Extended rift-associated volcanism in Ganis Chasma, Venus detected from Magellan radar emissivity

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Key Points

1. Radar anomalies in Ganis Chasma correspond to transient bright spots seen in infrared data
2. Low radar emissivity values suggest recent volcanic and tectonic activity on Venus
3. Lava flows in Ganis Chasma may have erupted over the last 30 years

Abstract

Exploration of Venus in the 1970–1990's revealed that the geology of Venus, the most Earth-like of the terrestrial planets, was decidedly un-Earth-like, with no plate tectonics, and no record of the first 80% of its history. A major outstanding question is whether Venus is still volcanically active today. We find that regions of Ganis Chasma have low radar emissivity values, due to low volumes of high dielectric minerals formed by surface – atmosphere weathering on the timescales of around 10s Ma. This confirms the presence of geologically recent volcanism in association with this major tectonic rift zone. The spatial correspondence of this emissivity signature with transient thermal anomalies suggests that Venus has been volcanically active at this site for at least the last few decades, a prediction that can be tested with space missions to Venus in the coming decade.

1 ONGOING VOLCANISM ON VENUS?

The diameter of Venus predicts that it should, like Earth, be volcanically active today (e.g., Head and Solomon, 1981). Radar images of the surface collected during the Magellan mission (1990–1994) did not identify any morphological evidence of recent volcanic activity. Nonetheless, recent and ongoing volcanic activity on Venus is suggested by other multiple independent lines of evidence. Pioneer Venus Orbiter (PV, 1978–1986) and Venus Express (VEx, 2005–2014) missions gathered more than 30 years of atmospheric measurements searching for evidence of possible volcanic eruptions. Both missions revealed fluctuations in sulfur dioxide (SO_2) that are possibly associated with volcanic outgassing on Venus (e.g., Esposito, 1984; Esposito et al., 1988; Marcq et al., 2012).

Idunn Mons, located in Imdr Regio, is considered to be among the most likely sites for active volcanism on the planet. Smrekar et al. (2010) detected some lava flows with relatively high thermal emission at 1 μm in the region. Infrared (1-μm) emissivity is derived from the Visual and Infrared Thermal Imaging Spectrometer (VIRTIS) images of the southern hemisphere returned by the VEx mission (Helbert et al., 2008; Mueller et al., 2008). In their work, Smrekar et al. (2010) suggest that the surface of Idunn Mons is made of young and unweathered lava flows, and hence, the volcano was possibly still active. Further investigations at Idunn reported that the young lava flows are more likely found on the eastern flanks of the volcano (D'Incecco et al., 2017), and contemporaneous to stratigraphically young structures associated with rifting at Olapa Chasma (D'Incecco et al., 2020; D'Incecco et al., 2021a; López et al., 2022). Other regions show a higher

1-µm emissivity relative to the surrounding plains, notably Hathor and Innini montes in Dione Regio (Mueller et al., 2008), Mielikki Mons, Shulamite and Shiwanokia coronae in Themis Regio (Stofan et al., 2009, 2016). The ages of the fresh, basaltic lava flows could not be well constrained since the timescale for chemical weathering in the near-surface environment of Venus is fundamentally unknown. The VIRTIS 1-µm emissivity data are expected to be controlled by the oxidation of ferrous iron in basalts to hematite, where the high thermal emission of Idunn Mons is thought to represent lower degrees of weathering; Smrekar et al. (2010) estimate that the volcano is 2.5 million years old or younger based on this reaction. More recent experimental studies under Venus conditions (Berger et al., 2019; Cutler et al., 2020; Filiberto et al., 2020; Filiberto et al., 2021; Teffeteller et al., 2022) suggest that chemical weathering occurs on much shorter time frames (i.e., in weeks to months).

Atla Regio is another region of prime interest for ongoing volcanic and tectonic activity, which is supported by analysis of the gravity and altimetry data returned by Magellan (Phillips, 1994; Smrekar, 1994; Stofan et al., 1995), and thermal anomalies observed by the Venus Monitoring Camera (VMC) during the VEx mission (Shalygin et al., 2012; 2015). Additionally, several analyses based on stratigraphic relationships between rift structures, lava flows, and crater features provide further evidence for the relative youthfulness of the region (e.g., Basilevsky, 1993; Basilevsky and Head, 2002a; 2002b; Brossier et al., 2021). VMC data collected during 2007–2009 did not reveal any signs of ongoing volcanic eruptions for the major volcanoes in Atla Regio (i.e., Maat, Ozza and Sapas montes) (Shalygin et al., 2012). However, they show several high (1-µm) emission spots with varying intensity over several days or months at different sites near Ganis Chasma (Shalygin et al., 2015). Ganis Chasma (or Ganiki Chasma) is a rift zone in Atla Regio centered at 192°E, 18°N, where recent activity was already suggested based on the superposition of rift structures on young impact deposits (Basilevsky, 1993). Shalygin et al. (2015) propose that these transient high emission spots are consistent with short-lived effusive activity, locally causing significant increases of surface temperatures.

2 MAGELLAN EMISSIVITY AS A CHRONOMETER

Radar emissivity can also be used as proxy to constrain the degree of weathering and therefore surface age of volcanic systems. Pioneer Venus and Magellan data show that many of Venus's highlands have distinctly elevated values of radar reflectivity (Masursky et al., 1980; Ford and

Pettengill, 1983) and thus low values of radar emissivity at their summits (Pettengill et al., 1992). These radar "anomalies" are ascribed to the presence of minerals with a high dielectric constant, **as** it is expected from theory that materials with high dielectric constants will enhance their radar reflectivity and lower their radar emissivity (Pettengill et al., 1992; Campbell, 1994). Several studies indicate that high dielectric minerals can be produced through chemical weathering reactions between the rocks and the near-surface atmosphere (e.g., Klose et al., 1992; Schaefer and Fegley, 2004; Treiman et al., 2016; Semprich et al., 2020 and references therein); if so, the reduction in radar emissivity can be associated with the formation of high dielectric minerals over time and thus can serve as a chronometer.

Brossier et al. (2020) reveal that most volcanoes and coronae on Venus are compatible with the presence of ferroelectric minerals in their rocks, particularly the tallest volcanoes on the planet (Maat and Ozza montes). Ferroelectric minerals (e.g., chlorapatite, perovskite oxides) are substances that undergo a phase transition when they reach a certain temperature, also called the Curie temperature, where its dielectric constant increases strongly. As the temperature rises above the Curie temperature (i.e., lower elevation on Venus), its dielectric constant gradually declines to normal values (Shepard et al., 1994; Treiman et al., 2016). Elevation and shape of the emissivity variations described in Brossier et al. (2020) indicate the presence of ferroelectrics with Curie temperatures of 693–731 K over a range of elevation between 6052.5 km and 6056.7 km. The varying "critical altitudes" reported in Klose et al. (1992) and seen by Brossier and colleagues could be due to diverse mineralogical compositions, or local differences in the atmospheric composition or temperature (Treiman et al., 2016). A more detailed investigation in Atla Regio (Brossier et al., 2021), shows that Maat and Ozza montes display multiple reductions in radar emissivity at different altitudes including, atypically, lowlands. These authors reported that these low emissivity signatures are found to correlate with individual lava flows, indicating that the excursions are controlled by variations in rock chemistry as opposed to the deposition of atmospheric precipitates.

Here we extract radar emissivity and elevation data collected during the Magellan mission (1990–1994) and examine the variation of emissivity with altitude for sites at Ganis Chasma identified as thermal anomalies in the VMC data (Shalygin et al., 2015), thus providing an independent constraint on surface age in the region. We believe that a detailed description of the radiophysical behaviors of these sites may help to retrieve, or at least constrain, their relative age and composition (as in Brossier et al., 2020; 2021; Brossier and Gilmore, 2021). The present paper is therefore organized as follows. We first locate and describe the changes in radar emissivity with altitude for the selected sites of interest in order to assess in detail the radiophysical signatures seen in Ganis Chasma (Section 4). This aims to determine whether the material measured in the region has the behavior consistent with that of known substances and we consider whether emissivity variations are related to rock age (Section 5).

3 DATA & METHODS

Our investigation uses radar datasets compiled during the Magellan mission (frequency = 2.4 GHz, λ = 12 cm). Morphological units are identified with the Cycle 1 left-looking Magellan Synthetic Aperture Radar (SAR) images (FMAPS) produced at a resolution of 75 m per pixel. The rift valley as well as the surrounding craters (e.g., Sitwell and Bashkirtseff craters), volcanoes (e.g., Yolkai-Estsan Mons) and tesserae were initially mapped in Ivanov and Head (2011) (Figure 1).

[Figure 1]

We derived altimetry and emissivity from the Magellan global topography data records (GTDR) and global emissivity data records (GEDR). Altimetry data have a spatial resolution ranging from ~10 km at periapsis (ca. 10˚N latitude) to ~20 km near the poles (ca. 90˚N and 70˚S) when the orbiting spacecraft was high above the planet. Emissivity data were collected while the spacecraft was operating in radiometer mode. The spatial resolution of the emissivity data varies from ~20 km near periapsis to ~80 km at high latitudes (Pettengill et al., 1991). Near-global mosaics are produced in the GTDR and GEDR data products that are publicly available through the USGS websites (https://planetarymaps.usgs.gov/mosaic). The two mosaics are resampled to a spatial resolution of 4.6 km per pixel (scale of 22.7 pixel per degree). Altimetry and emissivity data are extracted from these mosaics to produce scatterplots of the emissivity variation with altitude for each site of interest (e.g., Brossier et al., 2020; Brossier and Gilmore, 2021; Brossier et al., 2021), as in Klose et al. (1992). Elevation data are given in planetary radius with a mean value taken as 6051.8 km (Ford and Pettengill, 1992). Selection and extraction processes are done with the ArcGIS 10.6 (ESRI) software package, while the plots are produced with RStudio software. We also retrieve temperatures by correlation to the Vega 2 lander entry profile (Seiff, 1987; Lorenz et al., 2018; Brossier et al., 2020). Magellan datasets covering the study area, shapefiles (mapped units, and sites of interest), and extracted values (emissivity, altimetry and temperatures) are available through the online repository linked to this work (Brossier et al. 2022).

[Figure 2]

4 RESULTS

4.1 Study Sites

Our extraction is performed on the four sites studied with VMC data in Shalygin et al. (2015) (sites 1–4), and three other sites (sites 5–7) for comparison purposes. The main objective of this study is to use our methodology previously published (Brossier et al., 2020; 2021; Brossier and Gilmore et al., 2021) and to apply it on the exact same regions outlined in Shalygin et al. (2015), in order to have a direct comparison. Figure 1 displays the major morphological features in the region, while Figure 2 indicates the emissivity and elevation variations for each site. Sites 1 and 4 are located at the margins of the rift valley and replicate the boundaries of the strongest thermal anomalies identified by Shalygin et al. (2015). Both sites comprise outer flows and faulted walls of the rift valley. Sites 2 and 3 are also considered as areas of recent activity and correspond to high elevated and faulted walls of the rift valley. Among the new sites, 5 and 6 are morphologically similar to sites 2 and 3, and at similar high elevations. Site 7 corresponds to the extensive lava flows of Yolkai-Estsan Mons (hereafter called Yolkai for simplicity). This volcano has been heavily dissected by faults and is thus older than the rifting. Sitwell crater (32.8 km–diameter) has a parabolic ejecta deposit (parabola) that is superimposed on Ganis Chasma and may have undergone some rift-associated fracturing. This indicates possible continuation of rifting activity in this part of Ganis Chasma after the formation of the crater and its parabola (Basilevsky, 1993). Bashkirtseff crater (36.3 km–diameter) is another crater in the region that lacks a parabola and appears to be embayed by Yolkai lava flows.

4.2 Emissivity Excursions

Figure 3A shows elevation – emissivity plots obtained for the seven sites of interest. Because both composition and surface roughness can reduce emissivity, we distinguish emissivity values derived from the faulted walls of Ganis Chasma (red dots), from those related to flow materials at the edge of the rift (black dots) (see also Figure S1). Nonetheless, it is worth noting that this distinction may include some surrounding effects due to the difference in resolutions between SAR images (75 m per pixel) for the mapping of the lava flows and faulted walls, and the extraction of the elevation and emissivity data (4.6 km per pixel).

The magnitude of an emissivity excursion is defined by the percentage decrease between the minimum emissivity value observed in a region and the planetary average of ~0.85 (Pettengill et al., 1992). We observe different magnitudes and behaviors of the emissivity excursions: (1) a "strong" excursion is where emissivity shows a decrease of ~30% or more from the planetary average value, (2) a "subtle" excursion shows a decrease of 10–30%, or (3) no changes (≤ 10%) where emissivity is nearly constant with elevation. Figure 3B reports the magnitude of the emissivity excursions detected in each site and the corresponding altitude and temperature. Excursion magnitudes reported here are those of the lava flow units (black dots in Figure 3A), mitigating surface roughness effects. Sites 1–4 and site 6 have subtle declines in emissivity (11-21%) that reach minimum values of 0.672–0.753 at altitudes varying between 6054.2 km and 6055.8 km (701–716 K). Conversely, sites 5 and 7 have strong declines (~30%) to minimum values of 0.595–0.600 reached at 6056.2 km (697 K) and 6054.5 km (713 K), respectively. All values are summarized in Table 1 for all sites of interest.

[Figure 3]

[Table 1]

5 COMPOSITION & RELATIVE AGE

At each site, emissivity values gradually decline with increasing altitude from the lowlands (i.e., below 6053 km) to a given elevation (Figure 3A). This pattern of emissivity variations with altitude is consistent with ferroelectric behavior, characterized by a steady, gradual decline in radar emissivity with increasing elevation, then a sharp return to higher emissivity values at altitudes above 6056 km (around 700 K). Such a behavior is observed in Ovda Regio (Shepard et al., 1994; Treiman et al., 2016) and more globally in most volcanic edifices and tesserae on the planet (Brossier et al., 2020; Brossier and Gilmore, 2021). Ferroelectric minerals are known to be very conductive at a certain temperature, namely the mineral's Curie temperature (Tc). In Ganis Chasma, we see this behavior for site 1 (Figure 3A), and although the other sites do not reach elevations of 6056 km, the shape of the emissivity – elevation curve is similar to site 1 and other examples of ferroelectric behavior (Shepard et al., 1994; Treiman et al., 2016; Brossier et al., 2020; Brossier and Gilmore, 2021). In the ferroelectric model, the altitude (and temperature) of an emissivity excursion is a function of the composition, while its magnitude is a function of the volume of ferroelectric minerals (Shepard et al., 1994; Brossier et al., 2021). Chlorapatite and some perovskite oxides are good candidates, as their transition from ferro- to paraelectric occurs at temperatures found on the surface of Venus (690–735 K). The reader is referred to Brossier et al. (2021) for more details on the presence of ferroelectrics on Venus.

To use emissivity as a chronometer, we assume that the lava flows have a similar initial composition, and that the primary minerals in the flows are chemically weathered by the atmosphere over time to produce secondary minerals with high dielectric constants. In this model, sites with strong emissivity excursions occurring at high altitude (above 6053 km) are thought to have had enough time to produce the ferroelectric minerals responsible for the radar anomalies in the region via surface – atmosphere chemical weathering reactions. Conversely, sites with subtle or no emissivity excursions at high altitudes are considered to be young or possibly active since they have a lower volume of ferroelectric minerals. This model is supported by studies of other large volcanoes, such as Maat, Idunn and Otafuku montes, whose lava flows show subtle to low emissivity excursions that correlate with recent stratigraphic position (Brossier et al., 2020; 2021). In Ganis Chasma, the emissivity patterns imply that the youngest features are in sites 1, 3 and 4 (subtle to no emissivity excursions), while the oldest features are in sites 5 and 7 (strong emissivity excursions). This interpretation is in good agreement with the observations made using VMC images by Shalygin et al. (2015). In Ganis Chasma (and other rift valleys), rifting process may have an important role in faulting and creating freshly exposed rocks, and it would produce a signature similar to the newly erupted lava flows. Indeed, ferroelectric minerals would be formed or "triggered" in contact with the near-surface environment; thus, these detections may indicate the presence of very recent tectonic activity, in concert with the associate evidence for recent volcanism.

Shalygin et al. (2015) report that site 1 was the most prominent spot, followed by sites 2 and 3, while at site 4 it was uncertain if it was transient. It is worth noting that the IR-bright spots from VMC data are short-lived (only lasted a few days) and observed in 2008–2009. Conversely, our analysis displays older signatures from the early 1990's, leading to a 20 year-gap between the two observations. This suggests that site 2 has erupted since it was imaged by Magellan.

Overall, the sites have similar emissivity behaviors (variation with altitude) at comparable elevation ranges (Figure 3A), although they present different excursion magnitudes (i.e., different volume, age) and slightly different critical altitude (i.e., temperature, composition) (Figure 3B). Site 6 is very similar to site 2 in terms of emissivity excursions, although it is uncertain since the data points are more diffuse. Interestingly, site 7 has a distinct emissivity pattern, with a strong excursion at low elevation (below 6055 km) that resemble that of some volcanoes on Venus, such as Sekmet and Anala montes (Brossier et al., 2020). This slight variability in critical altitudes could be ascribed to slight differences in the ferroelectric composition, as discussed in Shepard et al. (1994) and Treiman et al. (2016). Shepard et al. (1994) demonstrate that minor change of the Pb abundance in a $(Pb,Ca)TiO_3$ perovskite can increase or decrease the Curie temperature (Rupprecht and Bell, 1964), and hence the critical altitude. For instance, a 1% change in the Pb abundance changes the Curie temperature by ~8 K, corresponding to a 1 km change in the transition altitude. Treiman et al. (2016) suggests that differences in anion composition (OH, F and Cl) or cation composition (substitution of Sr or rare Earth elements for Ca) in a $Ca_5(PO_4)_3(OH,F,Cl)$ apatite can also change the Curie temperature. More importantly, they state that chlorapatite is ferroelectric and thus the F:Cl ratio will control the Curie temperature where apatite with a larger F:Cl ratio would require higher temperatures (i.e., lower elevations) to exhibit a high dielectric constant (Rausch, 1976).

6 CONCLUSION

We show that the transient IR-bright spots detected in Shalygin et al. (2015) have radar emissivity values close to the planetary average (~0.85). Other regions in Ganis Chasma with similar morphology and elevation range have low emissivity values indicating the presence of minerals with a high dielectric constant (e.g., ferroelectrics), predicted to be produced by chemical weathering over time.

Sites 1, 3 and 4 are characterized by young materials, as they lack minerals with high dielectric constant (not yet produced). Sites 5 and 7 are characterized by older materials with a greater volume of these minerals. This is further supported for site 7 that has been dissected by the rift formation. All sites are consistent with the presence of ferroelectrics with subtle differences in the mineral composition (chlorapatite, or perovskite oxides). This is in agreement with the other volcanoes in Atla Regio, Maat and Ozza montes (Brossier et al., 2020; Brossier et al., 2021). The pattern of the radar emissivity in these regions is consistent with relatively young and unweathered materials. The transient IR-bright spots in these regions detected 20 years after Magellan, provide independent corroboration of active volcanism in Ganis Chasma since the 1990's.

As a possible site of current tectonic and volcanic activity, Atla Regio represents one important science target for the upcoming missions to Venus (see also D'Incecco et al., 2021b). Future missions will indubitably provide important clues about present-day activities on the planet (e.g., Glaze et al., 2018). NASA's Venus Emissivity, Radio Science, InSAR, Topography & Spectroscopy (VERITAS) mission (Smrekar et al., 2020) and ESA's EnVision mission (Ghail et al., 2012, 2020) will return complementary, critical datasets including improved topography, SAR imaging, gravity, and infrared spectroscopy. Additionally, NASA's Deep Atmosphere Venus Investigation of Noble gases, Chemistry, and Imaging (DAVINCI) mission (Garvin et al., 2022) will analyze gases typically extruded by active volcanoes (SO_2, CO_2, HCL, HF, and perhaps PH_3). Roscosmos' Venera-D mission (Senske et al., 2017; Zasova et al., 2019) will analyze the infrared (1-µm) emissivity at high resolution, while its lander will also provide in-situ geochemical measurements of the surface and subsurface composition.

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Data Availability Statement

All data used in our mapping and plotting procedures (e.g., Magellan images, shapefiles and extracted values for the sites of interest) can be found in the online repository linked to this study (Brossier et al. 2022). Magellan global datasets are also provided through the USGS website (https://planetarymaps.usgs.gov/mosaic), and described in Ford et al. (1993).

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Table(s)

Table 1 – Values for the seven sites of interest in Ganis Chasma. Sites 1–4 correspond to the locations of the VMC thermal anomalies indicated in Shalygin et al. (2015). Sites 5–7 are control areas with similar morphology and altitude range to sites 1–4.

Sites	Features	Lon. (°E)	Lat. (°N)	Area (km^2)	Minimum emissivity	Altitude (km)	Temp. (K) (*)	Excursion magnitude (%)
1	VMC anomaly	12.5	197.6	23300	0.718	6055.4	704.6	15.5
2	VMC anomaly	16.5	197.6	31100	0.672	6055.4	704.6	20.9
3	VMC anomaly	18.2	191.5	31700	0.718	6055.8	700.7	15.5
4	VMC anomaly	12.0	199.3	38200	0.753	6054.2	715.6	11.4
5	Control area	20.1	187.3	34440	0.595	6056.2	696.5	30.0
6	Control area	17.4	194.6	12200	0.684	6055.4	704.6	19.5
7	Control area	16.2	193.9	54600	0.600	6054.5	712.9	29.4

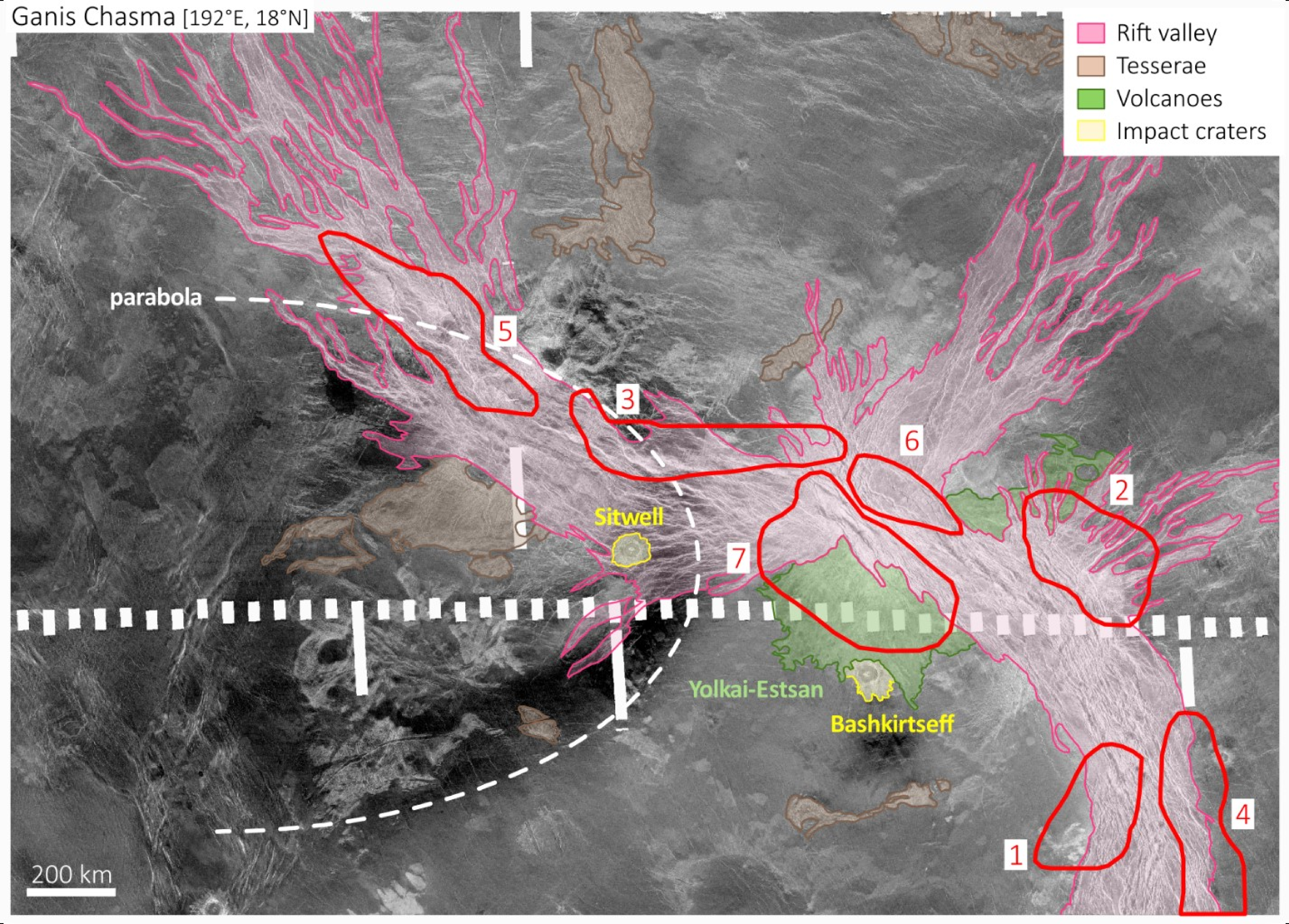
Notes: (*) Temperatures are derived from extrapolation of the Vega 2 lander data (Seiff, 1987; Lorenz et al., 2018) and reported in Brossier et al. (2020).

Figure Captions

Figure 1 – Ganis Chasma (192°E, 18°N) showing Magellan SAR image (gray scale) and the main morphologic features. The seven sites of interest are outlined in red. Morphologic features are mapped after Ivanov and Head (2011): Ganis Chasma (rift zone), Sitwell crater (with its parabola), Bashkirtseff crater, Yolkai-Estsan Mons, and surrounding tesserae. Maps (here and in Figure 2) have a simple cylindrical projection and north is up. Magellan images covering the study area and shapefiles (and auxiliary files) for the mapped units and sites of interest can be found in the online repository linked to this work (Brossier et al. 2022).

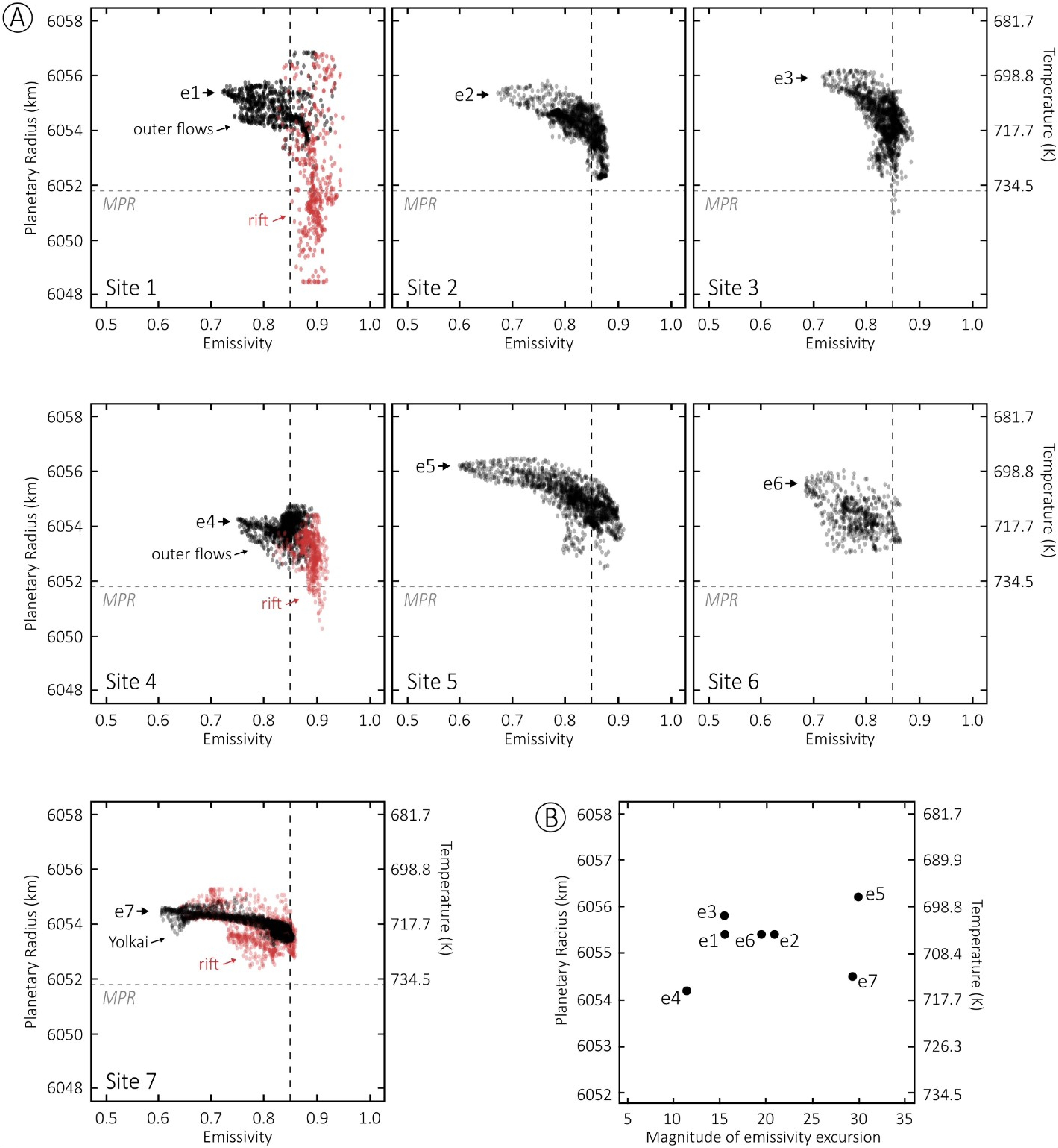
Figure 2 – Magellan radar emissivity and elevation overlapping SAR image (same as Figure 1) at Ganis Chasma (192°E, 18°N): (A) radar emissivity varies from low values in blue to high values in red, while (B) elevation varies from low elevations in teal to high elevation areas in brown.

Figure 3 – (A) Elevation vs emissivity plots obtained for the studied sites. Dashed lines in plots are mean global values of emissivity at 0.85 (vertical, black), and planetary radius at 6051.8 km (horizontal, gray). (B) Magnitude of emissivity excursions (percent change from global average value of 0.85) detected in each site vs. corresponding altitude and temperature. Temperatures are given by the Vega 2 lander data (Seiff, 1987; Lorenz et al., 2018; Brossier et al., 2020). Elevation (as planetary radius, in km) and emissivity values are reported as text files in the online repository (Brossier et al., 2022).



Ganis Chasma [192°E, 18°N]
Rift valley
Tesserae
Volcanoes
Impact craters
parabola
5
3
6
2
7
1
4
Sitwell
Yolkai-Estsan
Bashkirtseff
200 km


A
Radar Emissivity
0.595
0.970
parabola
Sitwell
Yolkai-Estsan
Bashkirtseff
1
2
3
4
5
6
7
200 km
B
Planetary Radius (km)
6048
6057
parabola
Sitwell
Yolkai-Estsan
Bashkirtseff
1
2
3
4
5
6
7
200 km


A
Site 1
Site 2
Site 3
Site 4
Site 5
Site 6
Site 7
e1
e2
e3
e4
e5
e6
e7
outer flows
rift
Yolkai
MPR
Planetary Radius (km)
Temperature (K)
Emissivity
B
Magnitude of emissivity excursion
e3
e1
e6
e2
e5
e7
e4